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Authors

Aab, A
Abreu, P
Aglietta, M
[et al.](#)

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A TARGETED SEARCH FOR POINT SOURCES OF EeV PHOTONS WITH THE
PIERRE AUGER OBSERVATORY

A. AAB⁶³, P. ABREU⁷⁰, M. AGLIETTA^{48,47}, I. AL SAMARAI²⁹, I.F.M. ALBUQUERQUE¹⁶, I. ALLEKOTTE¹, A. ALMELA^{8,11}, J. ALVAREZ CASTILLO⁶², J. ALVAREZ-MUÑOZ⁷⁹, G.A. ANASTASI³⁸, L. ANCHORDOQUI⁸³, B. ANDRADA⁸, S. ANDRINGA⁷⁰, C. ARAMO⁴⁵, F. ARQUEROS⁷⁷, N. ARSENE⁷³, H. ASOREY^{1,24}, P. ASSIS⁷⁰, J. AUBLIN²⁹, G. AVILA^{9,10}, A.M. BADESCU⁷⁴, A. BALACEANU⁷¹, R.J. BARREIRA LUZ⁷⁰, J.J. BEATTY⁸⁸, K.H. BECKER³¹, J.A. BELLIDO¹², C. BERAT³⁰, M.E. BERTAINA^{56,47}, X. BERTOU¹, P.L. BIERMANN¹⁰⁰¹, P. BILLOIR²⁹, J. BITEAU²⁸, S.G. BLAESS¹², A. BLANCO⁷⁰, J. BLAZEK²⁵, C. BLEVE^{50,43}, M. BOHÁČOVÁ²⁵, D. BONCIOLI^{40,1003}, C. BONIFAZI²², N. BORODAI⁶⁷, A.M. BOTTI^{8,33}, J. BRACK⁸², I. BRANCUS⁷¹, T. BRETZ³⁵, A. BRIDGEMAN³³, F.L. BRIECHLE³⁵, P. BUCHHOLZ³⁷, A. BUENO⁷⁸, S. BUITINK⁶³, M. BUSCEMI^{52,42}, K.S. CABALLERO-MORA⁶⁰, L. CACCIANIGA⁵³, A. CANCIO^{11,8}, F. CANFORA⁶³, L. CARAMETE⁷², R. CARUSO^{52,42}, A. CASTELLINA^{48,47}, G. CATALDI⁴³, L. CAZON⁷⁰, A.G. CHAVEZ⁶¹, J.A. CHINELLATO¹⁷, J. CHUDOBA²⁵, R.W. CLAY¹², R. COLALILLO^{54,45}, A. COLEMAN⁸⁹, L. COLLICA⁴⁷, M.R. COLUCCIA^{50,43}, R. CONCEIÇÃO⁷⁰, F. CONTRERAS^{9,10}, M.J. COOPER¹², S. COUTU⁸⁹, C.E. COVAULT⁸⁰, J. CRONIN⁹⁰, S. D'AMICO^{49,43}, B. DANIEL¹⁷, S. DASSO^{5,3}, K. DAUMILLER³³, B.R. DAWSON¹², R.M. DE ALMEIDA²³, S.J. DE JONG^{63,65}, G. DE MAURO⁶³, J.R.T. DE MELLO NETO²², I. DE MITRI^{50,43}, J. DE OLIVEIRA²³, V. DE SOUZA¹⁵, J. DEBATIN³³, O. DELIGNY²⁸, C. DI GIULIO^{55,46}, A. DI MATTEO^{51,41}, M.L. DÍAZ CASTRO¹⁷, F. DIOGO⁷⁰, C. DOBRIGKEIT¹⁷, J.C. D'OLIVO⁶², Q. DOROSTI³⁷, R.C. DOS ANJOS²¹, M.T. DOVA⁴, A. DUNDOVIC³⁶, J. EBR²⁵, R. ENGEL³³, M. ERDMANN³⁵, M. ERFANI³⁷, C.O. ESCOBAR¹⁰⁰⁵, J. ESPADANAL⁷⁰, A. ETCHEGOYEN^{8,11}, H. FALCKE^{63,66,65}, G. FARRAR⁸⁶, A.C. FAUTH¹⁷, N. FAZZINI¹⁰⁰⁵, B. FICK⁸⁵, J.M. FIGUEIRA⁸, A. FILIPČIĆ^{75,76}, O. FRATU⁷⁴, M.M. FREIRE⁶, T. FUJII⁹⁰, A. FUSTER^{8,11}, R. GAIOR²⁹, B. GARCÍA⁷, D. GARCIA-PINTO⁷⁷, F. GATÉ¹⁰⁰⁴, H. GEMMEKE³⁴, A. GHERGHEL-LASCU⁷¹, P.L. GHIA²⁸, U. GIACCARI²², M. GIAMMARCHI⁴⁴, M. GILLER⁶⁸, D. GLAS⁶⁹, C. GLASER³⁵, G. GOLUP¹, M. GÓMEZ BERISSO¹, P.F. GÓMEZ VITALE^{9,10}, N. GONZÁLEZ^{8,33}, A. GORGI^{48,47}, P. GORHAM⁹¹, A.F. GRILLO⁴⁰, T.D. GRUBB¹², F. GUARINO^{54,45}, G.P. GUEDES¹⁸, M.R. HAMPÉL⁸, P. HANSEN⁴, D. HARARI¹, T.A. HARRISON¹², J.L. HARTON⁸², A. HAUNGS³³, T. HEBBEKER³⁵, D. HECK³³, P. HEIMANN³⁷, A.E. HERVE³², G.C. HILL¹², C. HOJVAT¹⁰⁰⁵, E. HOLT^{33,8}, P. HOMOLA⁶⁷, J.R. HÖRANDEL^{63,65}, P. HORVATH²⁶, M. HRABOVSKÝ²⁶, T. HUEGE³³, J. HULSMAN^{8,33}, A. INSOLIA^{52,42}, P.G. ISAR⁷², I. JANDT³¹, S. JANSEN^{63,65}, J.A. JOHNSEN⁸¹, M. JOSEBACHUILI⁸, A. KÄÄPÄ³¹, O. KAMBEITZ³², K.H. KAMPERT³¹, I. KATKOV³², B. KEILHAUER³³, E. KEMP¹⁷, J. KEMP³⁵, R.M. KIECKHAFFER⁸⁵, H.O. KLAGES³³, M. KLEIFGES³⁴, J. KLEINFELLER⁹³, R. KRAUSE³⁵, N. KROHM³¹, D. KUEMPEL³⁵, G. KUKEC MEZEK⁷⁶, N. KUNKA³⁴, A. KUOTB AWAD³³, D. LAHURD⁸⁰, M. LAUSCHER³⁵, R. LEGUMINA⁶⁸, M.A. LEIGUI DE OLIVEIRA²⁰, A. LETESSIER-SELVON²⁹, I. LHENRY-YVON²⁸, K. LINK³², L. LOPES⁷⁰, R. LÓPEZ⁵⁷, A. LÓPEZ CASADO⁷⁹, Q. LUCE²⁸, A. LUCERO^{8,11}, M. MALACARI⁹⁰, M. MALLAMACI^{53,44}, D. MANDAT²⁵, P. MANTSCH¹⁰⁰⁵, A.G. MARIAZZI⁴, I.C. MARIS⁷⁸, G. MARSELLA^{50,43}, D. MARTELLO^{50,43}, H. MARTINEZ⁵⁸, O. MARTÍNEZ BRAVO⁵⁷, J.J. MASÍAS MEZA³, H.J. MATHES³³, S. MATHYS³¹, J. MATTHEWS⁸⁴, J.A.J. MATTHEWS⁹³, G. MATTHIAE^{55,46}, E. MAYOTTE³¹, P.O. MAZUR¹⁰⁰⁵, C. MEDINA⁸¹, G. MEDINA-TANCO⁶², D. MELO⁸, A. MENSHIKOV³⁴, M.I. MICHELETTI⁶, L. MIDDENDORF³⁵, I.A. MINAYA⁷⁷, L. MIRAMONTI^{53,44}, B. MITRICA⁷¹, D. MOCKLER³², S. MOLLERACH¹, F. MONTANET³⁰, C. MORELLO^{48,47}, M. MOSTAFÁ⁸⁹, A.L. MÜLLER^{8,33}, G. MÜLLER³⁵, M.A. MULLER^{17,19}, S. MÜLLER^{33,8}, R. MUSSA⁴⁷, I. NARANJO¹, L. NELLEN⁶², P.H. NGUYEN¹², M. NICULESCU-OGLINZANU⁷¹, M. NIECHCIOL³⁷, L. NIEMIETZ³¹, T. NIGGEMANN³⁵, D. NITZ⁸⁵, D. NOSEK²⁷, V. NOVOTNY²⁷, H. NOŽKA²⁶, L.A. NÚÑEZ²⁴, L. OCHIOI³⁷, F. OIKONOMOU⁸⁹, A. OLINTO²⁵, M. PALATKA²⁵, J. PALLOTTA², P. PAPPENBREER³¹, G. PARENTE⁷⁹, A. PARRA⁵⁷, T. PAUL^{87,83}, M. PECH²⁵, F. PEDREIRA⁷⁹, J. PEKALA⁶⁷, R. PELAYO⁵⁹, J. PEÑA-RODRIGUEZ²⁴, L. A. S. PEREIRA¹⁷, M. PERLÍN⁸, L. PERRONE^{50,43}, C. PETERS³⁵, S. PETRERA^{51,38,41}, J. PHUNTSOK⁸⁹, R. PIEGAIÁ³, T. PIEROG³³, P. PIERONI³, M. PIMENTA⁷⁰, V. PIRRONELLO^{52,42}, M. PLATINO⁸, M. PLUM³⁵, C. POROWSKI⁶⁷, R.R. PRADO¹⁵, P. PRIVITERA⁹⁰, M. PROUZA²⁵, E.J. QUEL², S. QUERCHFELD³¹, S. QUINN⁸⁰, R. RAMOS-POLLAN²⁴, J. RAUTENBERG³¹, D. RAVIGNANI⁸, B. REVENU¹⁰⁰⁴, J. RIDKY²⁵, M. RISSE³³, P. RISTORI², V. RIZI^{51,41}, W. RODRIGUES DE CARVALHO¹⁶, G. RODRIGUEZ FERNANDEZ^{55,46}, J. RODRIGUEZ ROJO⁹, D. ROGOZIN³³, M.J. RONCORONI⁸, M. ROTH³³, E. ROULET¹, A.C. ROVERO⁵, P. RUEHL³⁷, S.J. SAFFI¹², A. SAFTOIU⁷¹, F. SALAMIDA^{51,41}, H. SALAZAR⁵⁷, A. SALEH⁷⁶, F. SALESA GREUS⁸⁹, G. SALINA⁴⁶, F. SÁNCHEZ⁸, P. SANCHEZ-LUCAS⁷⁸, E.M. SANTOS¹⁶, E. SANTOS⁸, F. SARAZIN⁸¹, R. SARMENTO⁷⁰, C.A. SARMIENTO⁸, R. SATO⁹, M. SCHAUER³¹, V. SCHERINI⁴³, H. SCHIELER³³, M. SCHIMP³¹, D. SCHMIDT^{33,8}, O. SCHOLTEN^{64,1002}, P. SCHOVÁNEK²⁵, F.G. SCHRÖDER³³, A. SCHULZ³², J. SCHULZ⁶³, J. SCHUMACHER³⁵, S.J. SCIUTTO⁴, A. SEGRETO^{39,42}, M. SETTIMO²⁹, A. SHADKAM⁸⁴, R.C. SHELLARD¹³, G. SIGL³⁶, G. SILLI^{8,33}, O. SIMA⁷³, A. ŚMIAŁKOWSKI⁶⁸, R. ŠMÍDA³³, G.R. SNOW⁹², P. SOMMERS⁸⁹, S. SONNTAG³⁷, J. SOROKIN¹², R. SQUARTINI⁹, D. STANCA⁷¹, S. STANIĆ⁷⁶, J. STASIELAK⁶⁷, P. STASSI³⁰, F. STRAFELLA^{50,43}, F. SUAREZ^{8,11}, M. SUAREZ DURÁN²⁴, T. SUDHOLZ¹², T. SUOMIJÄRVI²⁸, A.D. SUPANITSKY⁵, J. SWAIN⁸⁷, Z. SZADKOWSKI⁶⁹, A. TABOADA³², O.A. TABORDA¹, A. TAPIA⁸, V.M. THEODORO¹⁷, C. TIMMERMANS^{65,63}, C.J. TODERO PEIXOTO¹⁴, L. TOMANKOVA³³, B. TOMÉ⁷⁰, G. TORRALBA ELIPE⁷⁹, P. TRAVNICEK²⁵, M. TRINI⁷⁶, R. ULRICH³³, M. UNGER³³, M. URBAN³⁵, J.F. VALDÉS GALICIA⁶², I. VALIÑO⁷⁹, L. VALORE^{54,45}, G. VAN AAR⁶³, P. VAN BODEGOM¹², A.M. VAN DEN BERG⁶⁴, A. VAN VLIET⁶³, E. VARELA⁵⁷, B. VARGAS CÁRDENAS⁶², G. VARNER⁹¹, J.R. VÁZQUEZ⁷⁷, R.A. VÁZQUEZ⁷⁹, D. VEBERIĆ³³, I.D. VERGARA QUISPE⁴, V. VERZI⁴⁶, J. VICHA²⁵, L. VILLASEÑOR⁶¹, S. VOROBIOV⁷⁶, H. WAHLBERG⁴, O. WAINBERG^{8,11}, D. WALZ³⁵, A.A. WATSON¹⁰⁰⁰, M. WEBER³⁴, A. WEINDL³³, L. WIENCKE⁸¹, H. WILCZYŃSKI⁶⁷, T. WINCHEN³¹, M. WIRTZ³⁵, D. WITTKOWSKI³¹, B. WUNDHEILER⁸, L. YANG⁷⁶, D. YELOS^{11,8}, A. YUSHKOV⁸, E. ZAS⁷⁹, D. ZAVRTANIK^{76,75}, M. ZAVRTANIK^{75,76}, A. ZEPEDA⁵⁸, B. ZIMMERMANN³⁴, M. ZIOLKOWSKI³⁷, Z. ZONG²⁸, F. ZUCCARELLO^{52,42}

¹Centro Atómico Bariloche and Instituto Balseiro (CNEA-UNCuyo-CONICET), Argentina²Centro de Investigaciones en Láseres y Aplicaciones, CITEDEF and CONICET, Argentina

- ³Departamento de Física and Departamento de Ciencias de la Atmósfera y los Océanos, FCEyN, Universidad de Buenos Aires, Argentina
- ⁴IFLP, Universidad Nacional de La Plata and CONICET, Argentina
- ⁵Instituto de Astronomía y Física del Espacio (IAFE, CONICET-UBA), Argentina
- ⁶Instituto de Física de Rosario (IFIR) – CONICET/U.N.R. and Facultad de Ciencias Bioquímicas y Farmacéuticas U.N.R., Argentina
- ⁷Instituto de Tecnologías en Detección y Astropartículas (CNEA, CONICET, UNSAM) and Universidad Tecnológica Nacional – Facultad Regional Mendoza (CONICET/CNEA), Argentina
- ⁸Instituto de Tecnologías en Detección y Astropartículas (CNEA, CONICET, UNSAM), Centro Atómico Constituyentes, Comisión Nacional de Energía Atómica, Argentina
- ⁹Observatorio Pierre Auger, Argentina
- ¹⁰Observatorio Pierre Auger and Comisión Nacional de Energía Atómica, Argentina
- ¹¹Universidad Tecnológica Nacional – Facultad Regional Buenos Aires, Argentina
- ¹²University of Adelaide, Australia
- ¹³Centro Brasileiro de Pesquisas Físicas (CBPF), Brazil
- ¹⁴Universidade de São Paulo, Escola de Engenharia de Lorena, Brazil
- ¹⁵Universidade de São Paulo, Inst. de Física de São Carlos, São Carlos, Brazil
- ¹⁶Universidade de São Paulo, Inst. de Física, São Paulo, Brazil
- ¹⁷Universidade Estadual de Campinas (UNICAMP), Brazil
- ¹⁸Universidade Estadual de Feira de Santana (UEFS), Brazil
- ¹⁹Universidade Federal de Pelotas, Brazil
- ²⁰Universidade Federal do ABC (UFABC), Brazil
- ²¹Universidade Federal do Paraná, Setor Palotina, Brazil
- ²²Universidade Federal do Rio de Janeiro (UFRJ), Instituto de Física, Brazil
- ²³Universidade Federal Fluminense, Brazil
- ²⁴Universidad Industrial de Santander, Colombia
- ²⁵Institute of Physics (FZU) of the Academy of Sciences of the Czech Republic, Czech Republic
- ²⁶Palacky University, RCPTM, Czech Republic
- ²⁷University Prague, Institute of Particle and Nuclear Physics, Czech Republic
- ²⁸Institut de Physique Nucléaire d’Orsay (IPNO), Université Paris-Sud, Univ. Paris/Saclay, CNRS-IN2P3, France, France
- ²⁹Laboratoire de Physique Nucléaire et de Hautes Energies (LPNHE), Universités Paris 6 et Paris 7, CNRS-IN2P3, France
- ³⁰Laboratoire de Physique Subatomique et de Cosmologie (LPSC), Université Grenoble-Alpes, CNRS/IN2P3, France
- ³¹Bergische Universität Wuppertal, Department of Physics, Germany
- ³²Karlsruhe Institute of Technology, Institut für Experimentelle Kernphysik (IEKP), Germany
- ³³Karlsruhe Institute of Technology, Institut für Kernphysik (IKP), Germany
- ³⁴Karlsruhe Institute of Technology, Institut für Prozessdatenverarbeitung und Elektronik (IPE), Germany
- ³⁵RWTH Aachen University, III. Physikalisches Institut A, Germany
- ³⁶Universität Hamburg, II. Institut für Theoretische Physik, Germany
- ³⁷Universität Siegen, Fachbereich 7 Physik – Experimentelle Teilchenphysik, Germany
- ³⁸Gran Sasso Science Institute (INFN), L’Aquila, Italy
- ³⁹INAF – Istituto di Astrofisica Spaziale e Fisica Cosmica di Palermo, Italy
- ⁴⁰INFN Laboratori Nazionali del Gran Sasso, Italy
- ⁴¹INFN, Gruppo Collegato dell’Aquila, Italy
- ⁴²INFN, Sezione di Catania, Italy
- ⁴³INFN, Sezione di Lecce, Italy
- ⁴⁴INFN, Sezione di Milano, Italy
- ⁴⁵INFN, Sezione di Napoli, Italy
- ⁴⁶INFN, Sezione di Roma “Tor Vergata“, Italy
- ⁴⁷INFN, Sezione di Torino, Italy
- ⁴⁸Osservatorio Astrofisico di Torino (INAF), Torino, Italy
- ⁴⁹Università del Salento, Dipartimento di Ingegneria, Italy
- ⁵⁰Università del Salento, Dipartimento di Matematica e Fisica “E. De Giorgi“, Italy
- ⁵¹Università dell’Aquila, Dipartimento di Scienze Fisiche e Chimiche, Italy

- ⁵²Università di Catania, Dipartimento di Fisica e Astronomia, Italy
- ⁵³Università di Milano, Dipartimento di Fisica, Italy
- ⁵⁴Università di Napoli “Federico II“, Dipartimento di Fisica “Ettore Pancini“, Italy
- ⁵⁵Università di Roma “Tor Vergata”, Dipartimento di Fisica, Italy
- ⁵⁶Università Torino, Dipartimento di Fisica, Italy
- ⁵⁷Benemérita Universidad Autónoma de Puebla (BUAP), México
- ⁵⁸Centro de Investigación y de Estudios Avanzados del IPN (CINVESTAV), México
- ⁵⁹Unidad Profesional Interdisciplinaria en Ingeniería y Tecnologías Avanzadas del Instituto Politécnico Nacional (UPIITA-IPN), México
- ⁶⁰Universidad Autónoma de Chiapas, México
- ⁶¹Universidad Michoacana de San Nicolás de Hidalgo, México
- ⁶²Universidad Nacional Autónoma de México, México
- ⁶³Institute for Mathematics, Astrophysics and Particle Physics (IMAPP), Radboud Universiteit, Nijmegen, Netherlands
- ⁶⁴KVI – Center for Advanced Radiation Technology, University of Groningen, Netherlands
- ⁶⁵Nationaal Instituut voor Kernfysica en Hoge Energie Fysica (NIKHEF), Netherlands
- ⁶⁶Stichting Astronomisch Onderzoek in Nederland (ASTRON), Dwingeloo, Netherlands
- ⁶⁷Institute of Nuclear Physics PAN, Poland
- ⁶⁸University of Łódź, Faculty of Astrophysics, Poland
- ⁶⁹University of Łódź, Faculty of High-Energy Astrophysics, Poland
- ⁷⁰Laboratório de Instrumentação e Física Experimental de Partículas – LIP and Instituto Superior Técnico – IST, Universidade de Lisboa – UL, Portugal
- ⁷¹“Horia Hulubei” National Institute for Physics and Nuclear Engineering, Romania
- ⁷²Institute of Space Science, Romania
- ⁷³University of Bucharest, Physics Department, Romania
- ⁷⁴University Politehnica of Bucharest, Romania
- ⁷⁵Experimental Particle Physics Department, J. Stefan Institute, Slovenia
- ⁷⁶Laboratory for Astroparticle Physics, University of Nova Gorica, Slovenia
- ⁷⁷Universidad Complutense de Madrid, Spain
- ⁷⁸Universidad de Granada and C.A.F.P.E., Spain
- ⁷⁹Universidad de Santiago de Compostela, Spain
- ⁸⁰Case Western Reserve University, USA
- ⁸¹Colorado School of Mines, USA
- ⁸²Colorado State University, USA
- ⁸³Department of Physics and Astronomy, Lehman College, City University of New York, USA
- ⁸⁴Louisiana State University, USA
- ⁸⁵Michigan Technological University, USA
- ⁸⁶New York University, USA
- ⁸⁷Northeastern University, USA
- ⁸⁸Ohio State University, USA
- ⁸⁹Pennsylvania State University, USA
- ⁹⁰University of Chicago, USA
- ⁹¹University of Hawaii, USA
- ⁹²University of Nebraska, USA
- ⁹³University of New Mexico, USA
-
- ¹⁰⁰⁰School of Physics and Astronomy, University of Leeds, Leeds, United Kingdom
- ¹⁰⁰¹Max-Planck-Institut für Radioastronomie, Bonn, Germany
- ¹⁰⁰²also at Vrije Universiteit Brussels, Brussels, Belgium
- ¹⁰⁰³now at Deutsches Elektronen-Synchrotron (DESY), Zeuthen, Germany
- ¹⁰⁰⁴SUBATECH, École des Mines de Nantes, CNRS-IN2P3, Université de Nantes
- ¹⁰⁰⁵Fermi National Accelerator Laboratory, USA

ABSTRACT

Simultaneous measurements of air showers with the fluorescence and surface detectors of the Pierre Auger Observatory allow a sensitive search for EeV photon point sources. Several Galactic and extragalactic candidate objects are grouped in classes to reduce the statistical penalty of many trials from that of a blind search and are analyzed for a significant excess above the background expectation. The presented search does not find any evidence for photon emission at candidate sources, and combined p -values for every class are reported. Particle and energy flux upper limits are given for selected candidate sources. These limits significantly constrain predictions of EeV proton emission models from non-transient Galactic and nearby extragalactic sources, as illustrated for the particular case of the Galactic center region.

Keywords: astroparticle physics — cosmic rays — methods: data analysis

1. INTRODUCTION

Ultra-high energy (UHE) photons with energies around 1 EeV ($1 \text{ EeV} = 10^{18} \text{ eV}$) and above have not yet been identified (see [Bleve & the Pierre Auger Collaboration \(2015\)](#) and references therein). At these high energies photons are produced primarily by π^0 decays, implying the existence of hadrons (that cause the production of π^0 mesons) with energies typically 10 times higher than the secondary photon ([Risse & Homola 2007](#)). At energies of about 5 EeV, around the “ankle” of the energy spectrum ([Abraham et al. 2010](#); [Schulz & the Pierre Auger Collaboration 2013](#)), several experiments, including the Pierre Auger Observatory, HiRes, and Telescope Array, have all found their measurements to be consistent with the existence of a light component among the cosmic rays ([Abreu et al. 2012a](#); [Aab et al. 2014a,b, 2016](#); [Abbasi et al. 2010](#); [Jui & the Telescope Array Collaboration 2012](#)). If these protons were to interact in the vicinity of their sources they can produce photons by pion photoproduction or inelastic nuclear collisions. Since photons are not deflected by magnetic fields, the experimental signature would then be an accumulation of photon-like events from a particular celestial direction.

Assuming that the energy spectra of measured TeV γ sources ([Hinton & Hofmann 2009](#); [Abramowski et al. 2011](#)) extend to EeV energies, it is plausible that photon and neutron fluxes are also detectable in the EeV energy range. Sources producing particle fluxes according to an E^{-2} energy spectrum inject equal energy into each decade. A measured energy flux of $1 \text{ eV cm}^{-2} \text{ s}^{-1}$ in the TeV decade would result in the same energy flux in the EeV decade if the spectrum continues to such high energies and energy losses en route to Earth are negligible (see Section 3). Furthermore, the H.E.S.S. collaboration measured a TeV gamma ray spectrum from the Galactic center region without any observation of a cutoff or a spectral break up to tens of TeV, implying that our Galaxy hosts petaelectronvolt accelerators

called “PeVatrons” ([Abramowski et al. 2016](#)). If these photons are produced in hadronic processes they are necessarily accompanied by neutrons produced in charge exchange interactions of protons. The ratio between photon and neutron emissivities from p - p collisions at the same pivot energy depends primarily on the spectral index of the proton source and it is shown in [Crocker et al. \(2005\)](#) that for spectral indices $\Gamma_p \lesssim 2.4$ photon emissivities dominate, assuming a continuation of the parent proton spectrum well beyond the pivot energy. Several experiments, including the Pierre Auger Observatory, Telescope Array, IceCube, and KASCADE, searched for an indication of neutron fluxes above the PeV energy range from specific source directions, but no significant excess or correlation with catalogs could be found (e.g. [Aab et al. \(2014d\)](#); [Abbasi et al. \(2015\)](#); [Antoni et al. \(2004\)](#); [Aartsen et al. \(2016\)](#)).

This paper reports on a targeted search for photon point sources at EeV energies and complements previous neutron searches. The search for a photon flux, as opposed to a neutron flux, has a more direct connection to TeV measurements where the messengers are photons. A neutron flux is limited by decay of the neutrons with a mean path length of $9.2 \times E \text{ [EeV] kpc}$, requiring an energy of at least 1 EeV to observe the Galactic center region. In this paper we apply a lower energy threshold of $10^{17.3} \text{ eV}$ using events measured by the air fluorescence detector (FD) as well as the surface detector (SD) of the Pierre Auger Observatory (see Section 2). This choice provides high event statistics despite the reduced duty cycle of the FD. The sensitivity to photon point sources is increased by reducing the hadronic background contribution using mass-sensitive observables. In the case of neutron-induced air showers that is not possible, since they are indistinguishable from proton primaries.

In a previous paper the Pierre Auger Collaboration published the directional search for EeV photon point sources from any direction in the exposed sky (blind search; [Aab et al. \(2014c\)](#)). That analysis did not find a statistical evidence for any photon flux. The detected

small p -values are within the expectation given 526,200 target centers. To reduce the statistical penalty of many trials from that of a blind search, this analysis focuses on just 12 target sets, each set being a class of possible sources of high-energy photons (Section 3). The candidate sources all lie in the declination range -85° and $+20^\circ$. Targets in each class are combined in a “stacked analysis,” assuming that most or all candidate sources in a target set are emitting photons resulting in a more significant combined signal compared to that of a single target (Section 4). The results of this analysis, including particle and energy flux upper limits of selected target directions, are given in Section 5. This study uses the same methods for hadron reduction and calculation of upper limits that were explained in the preceding paper.

2. DATA SET

Air showers induced by UHE cosmic rays detected with the Pierre Auger Observatory (The Pierre Auger Collaboration 2015) are used in this analysis. The observatory is located in Argentina near Malargüe and is centered at latitude 35.2° S and longitude 69.5° W at a mean altitude of 1400 m above sea level. A SD of 1660 water-Cherenkov particle detectors covering an area of 3000 km^2 on a triangular grid with 1.5 km spacing observes electrons, muons, and photons at the ground with a duty cycle of nearly 100%. The area is overlooked by 27 fluorescence telescopes operating on dark nights with a duty cycle of $\sim 15\%$. Events recorded between 2005 January and 2013 December in hybrid mode, i.e., recorded by both the FD and SD, are used in this analysis. The selection criteria are the same as in the previous blind search paper (Aab et al. 2014c), but additional accumulated data increase the statistics by 28% to 308,676 recorded events in the present study. The energy range is between $10^{17.3} \text{ eV}$ and $10^{18.5} \text{ eV}$ to take advantage of high statistics at low energies and to avoid additional shower development processes at the highest energies (Risse & Homola 2007). The average angular resolution of the final dataset is 0.7° .

3. TARGET SET

The detectable source distance is limited by interactions of UHE photons with low-energy background photons in pair or double-pair production processes. The attenuation length, i.e., the distance at which the survival probability has dropped to $e^{-1} \simeq 36.8\%$, depends on the energy of the UHE photon. The expected attenuation length for photons in interactions with the cosmic microwave background (dominating) and with radio (Protheroe & Biermann 1996) and infrared (Kneiske et al. 2004) photon fields is shown as the solid black line in Figure 1. In the energy range

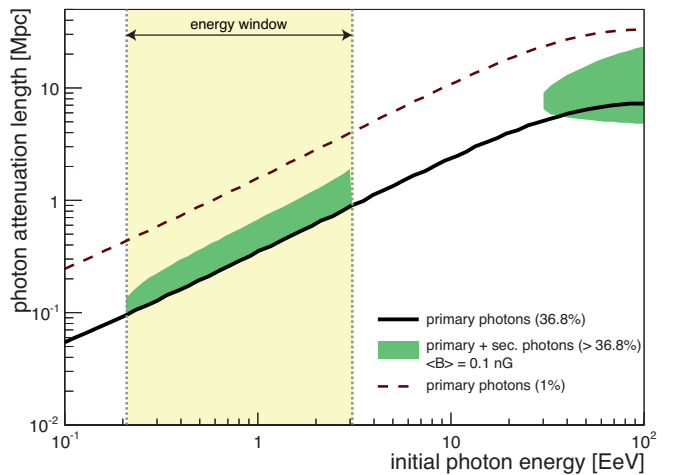


Figure 1. Photon attenuation length as a function of the initial primary energy. The thick black line indicates the attenuation length (survival probability $e^{-1} \simeq 36.8\%$) and the dashed line indicates a reduced survival probability of 1%. The energy range of this paper between $10^{17.3} \text{ eV}$ and $10^{18.5} \text{ eV}$ is indicated by the yellow shaded region and the vertical dotted lines. The expected increase of the observable distance by including secondary photons (detected in the energy range of this paper and with less than 1° deflection with respect to the primary photon) is shown as the green area using an average magnetic field strength of 0.1 nG.

of the dataset, the attenuation length varies between 90 kpc at $10^{17.3} \text{ eV}$ and 900 kpc at $10^{18.5} \text{ eV}$ encompassing Galactic and nearby extragalactic sources. Requiring only a survival probability of 1% the attenuation length increases, as indicated by the dashed line in Figure 1, extending to a distance of a few Mpc at the highest energies considered and hence also including the nearest active galactic nucleus Centaurus A, which lies at less than 4 Mpc. It should be noted that the observable distance may increase further if including also secondary photons, i.e., UHE photons converted back from electrons via inverse Compton scattering within the electromagnetic cascading process taking place outside the Galaxy. However, to be useful for point source searches, these photons must be in the energy range and have parent electrons that have not been deflected more than 1° in ambient magnetic fields (see Section 4). The expected increase of the maximum observable distance in a Kolmogorov-type turbulent field with a mean magnetic field strength $\langle B \rangle = 0.1 \text{ nG}$ is shown as the green shaded area in Figure 1, applying a three-dimensional CRPropa 3 simulation (Alves Batista et al. 2016). Note that even primary photon energies outside the energy window above $10^{18.5} \text{ eV}$ become visible if there is sufficient time to form the electromagnetic cascade. However, these results are very sensitive to the uncertain extragalactic magnetic field assumption, e.g., the maximum observable distance would drop to the values corresponding to the primary photon line if the mean mag-

Table 1. Combined Unweighted probabilities \mathcal{P} and Weighted Probabilities \mathcal{P}_w for the 12 Target Sets.

Note. In addition, information on the most significant target from each target set is given. The number of observed (Obs) and expected (Exp) events and the corresponding exposure are shown. The numbers in brackets in the observed number of events column indicate the number of events needed for a 3σ observation unpenalized and penalized (*). Upper limits (UL) are computed at 95% confidence level. The last two columns indicate the p -value unpenalized (p) and penalized (p^*). Due to the discrete distribution of p -values arising in isotropic simulations, \mathcal{P} can differ from p in the sets that contain only a single target.

Class	No.	\mathcal{P}_w	\mathcal{P}	R.A.	Decl.	Obs	Exp	Exposure	Flux UL	E -flux UL	p	p^*
				[$^\circ$]	[$^\circ$]			[$\text{km}^2 \text{ yr}$]	[$\text{km}^{-2} \text{ yr}^{-1}$]	[$\text{eV cm}^{-2} \text{ s}^{-1}$]		
msec PSRs	67	0.57	0.14	286.4	4.0	5 (7,9*)	1.433	236.1	0.043	0.077	0.010	0.476
γ -ray PSRs	75	0.97	0.98	312.8	-8.5	6 (8,10*)	1.857	248.1	0.045	0.080	0.007	0.431
LMXB	87	0.13	0.74	258.1	-40.8	6 (8,11*)	2.144	233.9	0.046	0.083	0.014	0.718
HMXB	48	0.33	0.84	285.9	-3.2	4 (7,9*)	1.460	235.2	0.036	0.066	0.040	0.856
H.E.S.S. PWN	17	0.92	0.90	266.8	-28.2	4 (8,10*)	2.045	211.4	0.038	0.068	0.104	0.845
H.E.S.S. other	16	0.12	0.52	258.3	-39.8	5 (8,10*)	2.103	233.3	0.040	0.072	0.042	0.493
H.E.S.S. UNID	20	0.79	0.45	257.1	-41.1	6 (8,10*)	2.142	239.2	0.045	0.081	0.014	0.251
Microquasars	13	0.29	0.48	267.0	-28.1	5 (8,10*)	2.044	211.4	0.045	0.080	0.037	0.391
Magnetars	16	0.30	0.89	257.2	-40.1	4 (8,10*)	2.122	253.8	0.031	0.056	0.115	0.858
Gal. Center	1	0.59	0.59	266.4	-29.0	2 (8,8*)	2.048	218.9	0.024	0.044	0.471	0.471
LMC	3	0.52	0.62	84.4	-69.2	2 (8,9*)	2.015	180.3	0.030	0.053	0.463	0.845
Cen A	1	0.31	0.31	201.4	-43.0	3 (8,8*)	1.948	214.1	0.031	0.056	0.221	0.221

netic field were $\langle B \rangle > 1$ nG, since in this case the electrons would be largely deflected.

Since there is a close connection between hadronic production processes for photons and neutrons, any candidate source of neutrons is also a candidate source of photons. As a consequence this analysis adopts the Galactic point source target sets defined in [Aab et al. \(2014d\)](#) but adds the new H.E.S.S. unidentified sources reported in [Deil et al. \(2015\)](#). The Galactic source classes are millisecond pulsars (msec PSRs), γ -ray pulsars (γ -ray PSRs), low-mass and high-mass X -ray binaries (LMXBs and HMXBs), H.E.S.S. Pulsar Wind Nebulae (PWNe), other H.E.S.S. identified and unidentified sources, microquasars, magnetars, and the Galactic center. To retain independent target sets a candidate source that appears in two or more sets is kept only in the most exclusive set. Because the maximum observable distance of EeV photons is greater than that for EeV neutrons, two additional extragalactic target sets are included in this analysis. One set consists of three powerful gamma-ray emitters in the Large Magellanic Cloud (LMC) at a distance of ~ 50 kpc ([Abramowski et al. 2015](#)). The core region of Centaurus A (Cen A) is, by itself, the second extragalactic target set. The 12 source classes collectively include 364 individual candidate source directions.

4. ANALYSIS METHOD

To reduce the contamination of hadronic background events, only air showers similar to the photon expectation are selected using the multivariate method of

Boosted Decision Trees ([Breiman et al. 1984](#); [Schapire 1990](#)) trained with Monte Carlo simulations of showers produced by photon and proton primaries. For a fixed primary energy, photon induced air showers have, on average, a delayed shower development and fewer muons (mostly electromagnetic component) compared to hadron-induced showers. As in the previous photon search paper, five different mass-sensitive observables are used: the depth of shower maximum X_{max} (from FD, being sensitive to delayed shower development), reduced χ^2 and normalized energy of the Greisen fit to the longitudinal profile (from FD, being sensitive to the electromagnetic component), S_b -parameter ([Ros et al. 2011](#)) (from SD, being sensitive to the slope of the lateral distribution of the shower, and hence to the muonic content), and the ratio of the early arriving to the late arriving integrated signal in the detector with the strongest signal (from SD, being sensitive to the muonic component and to the delayed shower development).

The optimized cut in the multivariate output distribution for a specific candidate source direction i depends on the expected number of isotropic background events b_i . This number is calculated by applying the scrambling technique ([Cassiday et al. 1990](#)), and naturally takes into account detector efficiencies and aperture features by assigning arrival times and arrival directions, binned for each telescope, randomly from measured events. This procedure is repeated 5000 times and the mean number of arrival directions within a target is then used as the expected isotropic background count.

For each target direction we use a top-hat counting region of 1° . Details of this multivariate cut selection and counting procedure are given in [Aab et al. \(2014c\)](#). Averaging over all 364 target directions, the multivariate cut is expected to retain 81.4% of photons while rejecting 95.2% of background hadrons. After applying the cut, the total number of recorded events from all of the targets is reduced from 11,180 to 474.

Each target set is tested with and without statistical weights. The weight w_i is assigned to each target i in a target set proportional to the measured electromagnetic flux f_i in the catalog and proportional to the directional photon exposure ϵ_i of the Pierre Auger Observatory based on [Settimo & The Pierre Auger Collaboration \(2012\)](#). Relative attenuation differences from candidate sources of the same class can be neglected given an interaction length between 90 and 900 kpc of primary photons in the energy range considered (see [Figure 1](#)). The sum of weights in each set is normalized to 1 (see [Aab et al. \(2014d\)](#)):

$$w_i = \frac{f_i \cdot \epsilon_i}{\sum_i f_i \cdot \epsilon_i}. \quad (1)$$

A p -value p_i is assigned to each candidate source of a target set as follows. The p -value for the target i is defined by $p_i \equiv [\text{Poisson}(n_i, b_i) + \text{Poisson}(n_i + 1, b_i)]/2$, where $\text{Poisson}(n_i, b_i)$ is the probability of getting n_i or more arrival directions in the target when the observed value is n_i , and the expected number from the background is b_i . Averaging the values for n and $n+1$ avoids a bias toward low or high p -values for pure background fluctuations.

The combined weighted probability \mathcal{P}_w is the fraction of isotropic simulations yielding a weighted product $\prod_i p_{i,\text{iso}}^{w_i}$ that is not greater than the measured weighted product $\prod_i p_i^{w_i}$:

$$\mathcal{P}_w = \text{Prob} \left(\prod_i p_{i,\text{iso}}^{w_i} \leq \prod_i p_i^{w_i} \right), \quad (2)$$

where $p_{i,\text{iso}}$ denotes the p -value of target i in an isotropic simulation. The combined unweighted probability \mathcal{P} is given by the same formula with $w_i = 1$ for all targets (see [Aab et al. \(2014d\)](#)).

5. RESULTS

The results for the combined analysis for each of the 12 target sets are shown in [Table 1](#), along with detailed information about the target that has the smallest p -value in each set. In addition to the direction of the candidate source, the measured and expected numbers of events within an opening angle of 1° are given along with the required number of events for a 3σ observation. In the last two columns are the minimum p -value of the target set (p) and the penalized p -value $p^* = 1 - (1 - p)^N$,

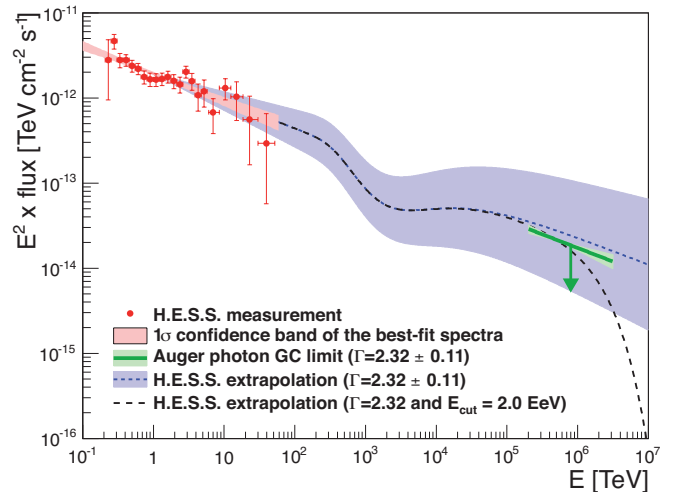


Figure 2. Photon flux as a function of energy from the Galactic center region. Measured data by H.E.S.S. are indicated, as well as the extrapolated photon flux at Earth in the EeV range, given the quoted spectral indices ([Abramowski et al. \(2016\)](#)); conservatively the extrapolation does not take into account the increase of the p - p cross-section toward higher energies). The Auger limit is indicated by a green line. A variation of the assumed spectral index by ± 0.11 according to systematics of the H.E.S.S. measurement is denoted by the light green and blue band. A spectral index with cutoff energy $E_{\text{cut}} = 2.0 \cdot 10^6$ TeV is indicated as well.

which is the chance probability that one or more of the N candidate sources in the target set would have a p -value less than p if the N p -values were randomly sampled from the uniform probability distribution.

No combined p -value (\mathcal{P} or \mathcal{P}_w) nor any individual target p -value has a statistical significance as great as 3σ . Upper limits are therefore derived for the flux from the target of smallest p -value in each target set assuming an E^{-2} photon spectrum and they are indicated in [Table 1](#). Upper limits on the photon flux from a point source i are calculated as $f_i^{95\%} = n_i^{\text{Zech}} / (n_{\text{inc}} \cdot \epsilon_i)$, where n_i^{Zech} is the upper limit, at the 95% confidence level, on the number of photons using Zech's definition ([Zech 1989](#)), $n_{\text{inc}} = 0.9$ is the expected signal fraction within the search window, and ϵ_i is the directional photon exposure.

Various sources of systematic uncertainties have been investigated in [Aab et al. \(2014c\)](#). The main contribution arises from the unknown photon spectral index due to the associated change in the directional photon exposure. Differences in the particle flux upper limit of -34% and $+51\%$ have been estimated when changing the photon spectral index from 2.0 to 1.5 or 2.5, respectively. Considering the background rejection, differences in the hadronic interaction models change the particle flux upper limits by, on average, -9% when using EPOS-LHC ([Pierog et al. 2013](#)) for proton simulations instead of QGSJET-01c ([Kalmykov & Ostapchenko 1989](#)).

In the following, the limit on the Galactic center

is examined in more detail. This is of particular interest, as the H.E.S.S. collaboration recently reported an indication of the acceleration of PeV protons from this region (Abramowski et al. 2016). H.E.S.S. measured the diffuse γ -ray emission following a power-law spectrum according to $dN/dE = \Phi_1 E^{-\Gamma}$ with a spectral index of $\Gamma = 2.32 \pm 0.05_{\text{stat}} \pm 0.11_{\text{syst}}$ and flux normalization $\Phi_1 = (1.92 \pm 0.08_{\text{stat}} \pm 0.28_{\text{syst}}) \times 10^{-12} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ without a cutoff or break up to tens of TeV (Abramowski et al. 2016). Since the results in Table 1 are based on a photon spectral index of $\Gamma = 2$ the limit is recalculated assuming $\Gamma = 2.32$ resulting in a particle flux upper limit at the 95% confidence level of the Galactic center region of $J_{\text{int}}^{\text{UL}} = 0.034 \text{ km}^{-2} \text{ yr}^{-1}$. As can be seen in Figure 2, the current photon flux upper limit can severely constrain the allowed parameter space for a flux continuation to EeV energies. This extrapolation takes into account interactions with the cosmic microwave background (dominating) and with radio (Protheroe & Biermann 1996) and infrared (Gilmore et al. 2012) photon fields. Furthermore, assuming a power law with an exponential cutoff of the form $dN/dE = \Phi_1 E^{-\Gamma} \times \exp(-E/E_{\text{cut}})$ an upper limit of the cutoff energy $E_{\text{cut}} = 2.0 \text{ EeV}$ can be placed by setting $\int_{E_1}^{E_2} \Phi_1 E^{-\Gamma} \times \exp(-E/E_{\text{cut}}) dE = J_{\text{int}}^{\text{UL}}$ with $E_1 = 10^{17.3} \text{ eV}$ and $E_2 = 10^{18.5} \text{ eV}$ and solving for E_{cut} . The corresponding cutoff spectrum is also given by the dashed line in Figure 2.

6. DISCUSSION

No target class reveals compelling evidence for photon-emitting sources in the EeV energy regime. For the 12 sets, the minimum combined weighted probability \mathcal{P}_w is 0.12. With 12 trials, one expects a \mathcal{P}_w -value at least that small to occur by chance, with 78% confidence. The minimum unweighted \mathcal{P} -value, 0.14, is similarly not statistically significant. There is also no evidence for one outstanding target in any target set. The minimum penalized p -value p^* in the 12 sets is 0.221. The null result holds also against the hypothesis that only a subset of some target class contributes a photon flux. This has been tested by calculating combined \mathcal{P} -values scanning only over the most significant, i.e., the smallest p -value, targets in the catalog.

The results presented in this paper complement previous results published by the Pierre Auger Collaboration searching for neutrons at higher energies using SD data (Abreu et al. 2012b; Aab et al. 2014d), and photons using hybrid data (Aab et al. 2014c), by restricting the photon search to 12 predefined target classes. Flux upper limits from photon point sources constrain the continuation of measured TeV fluxes to EeV energies, as shown for the particular case of the Galactic center introducing an upper limit of the cutoff photon energy

of $E_{\text{cut}} = 2.0 \text{ EeV}$.

The discovery of photon fluxes from any target set or individual targets in this study would have proved that EeV protons are being accelerated at discrete sources within the Galaxy or its neighborhood. The null results reported here leave open the possibility that EeV protons, as observed on Earth, are of extragalactic origin. Some support for that hypothesis was noted in the large-scale anisotropy analysis of Auger data (Abreu et al. 2013). It is important to note, however, that the absence of detectable photon fluxes, as reported here, does not exclude the production of EeV protons within the Galaxy. The derived flux limits are time-averaged values. EeV photons might be produced in transient sources, such as gamma-ray bursts or supernovae, or aligned in jets not pointing to us. An alternative explanation is that EeV protons escape from a source more freely than protons that produce TeV photon fluxes, and the production of EeV photons is thereby too meager to be detectable in the present study.

With the detector upgrade AugerPrime (Engel & the Pierre Auger Collaboration 2016; The Pierre Auger Collaboration 2016) the photon-hadron separation will be further improved, allowing an increased sensitivity for photon point sources.

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